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SMASH model and atmospheric neutrino mass splitting

C.R. Das\textsuperscript{1}, Katri Huitu\textsuperscript{2}, Timo Kärkkäinen\textsuperscript{3}

\textsuperscript{1}Bogoliubov Laboratory of Theoretical Physics, Joint Institute for Nuclear Research, Joliot-Curie 6, Dubna 141980, Moscow region, Russian Federation
\textsuperscript{2}Helsinki Institute of Physics, P.O.Box 64 (Gustaf Hällströminkatu 2), FI-00014 University of Helsinki, Finland
\textsuperscript{3}Department of Physics, P.O. Box 64 (Gustaf Hällströminkatu 2), FI-00014 University of Helsinki, Finland

E-mail: \textsuperscript{1}das@theor.jinr.ru, \textsuperscript{2}katri.huitu@helsinki.fi, \textsuperscript{3}timo.j.karkkainen@helsinki.fi

Abstract. Five fundamental problems - neutrino mass, baryogenesis, dark matter, inflation, strong CP problem - are solved at one stroke in a model, dubbed as “SM-A-S-H” (Standard Model-Axion-Seesaw-Higgs portal inflation) by Andreas Ringwald et. al. The Standard Model (SM) particle content is extended by three right-handed SM-singlet neutrinos $N_i$, a vector-like color triplet quark $Q$, a complex SM-singlet scalar field $\sigma$ that stabilises the Higgs potential, all of them being charged under Peccei-Quinn (PQ) $U(1)$ symmetry, the vacuum expectation value $v_\sigma \sim 10^{11}$ GeV breaks the lepton number and the Peccei-Quinn symmetry simultaneously. We found that numerically SMASH model not only solves five fundamental problems but also the sixth problem “Vacuum Metastability” through the extended scalar sector and can predict approximately correct atmospheric neutrino mass splitting around 0.05 eV and the solar neutrino mass splitting around 0.009 eV.

1. Introduction

After the discovery of the Standard Model (SM) Higgs boson [1, 2, 3], every elementary particle of the SM has been confirmed to exist. Even though the last forty years have been a spectacular triumph for the SM, the mass of the Higgs boson ($m_h = 125.09 \pm 0.32$ GeV) poses a serious problem for the SM. It is well-known that the SM Higgs potential is metastable [4], as the sign of the quartic coupling, $\lambda_H$ turns negative at instability scale $\Lambda_{IS} \sim 10^{11}$ GeV. The largest uncertainties of SM vacuum stability are driven by top quark pole mass and the mass of SM Higgs boson. Present experimental data is in significant tension with the stability hypothesis, making it more likely that the universe is in a metastable vacuum state. The expected lifetime of vacuum decay to a stable vacuum is extraordinarily long, and it is unlikely to affect the evolution of the universe. However, it is unclear why the vacuum state entered to metastable or unstable vacuum, to begin with during the early universe. On this post-SM era, the emergence of vacuum stability problem (among many others) forces the particle theorists to expand the SM in such a way that the $\lambda_H$ will stay positive during the running all the way up to the Planck scale, $M_P = 1.22 \times 10^{19}$ GeV.

It is possible that at or below the instability scale heavy degrees of freedom originating from a theory beyond the SM start to alter the running of the SM parameters of renormalization.
Consider an energy scale below $2.1$. Threshold correction which has eigenvalues where light degrees of freedom are ignored and $\lg$ denotes the Briggsian logarithm.

It turns out that the quartic coupling we measure has an additional term: the low-energy Higgs potential should match the SM Higgs potential:

$$V(H,\sigma) = \lambda_H \left( H^\dagger H - \frac{v^2}{2} \right)^2 + \lambda_\sigma \left( |\sigma|^2 - \frac{v_\sigma^2}{2} \right)^2 + 2\lambda_{H\sigma} \left( H^\dagger H - \frac{v^2}{2} \right) \left( |\sigma|^2 - \frac{v_\sigma^2}{2} \right).$$

Defining $\phi_1 = H$ and $\phi_2 = \sigma$, scalar mass matrix of this potential is

$$\left( M_{ij} \right) = \frac{1}{2} \left. \frac{\partial^2 V}{\partial \phi_i \partial \phi_j} \right|_{H=v/\sqrt{2},\sigma=v_\sigma/\sqrt{2}} = \begin{pmatrix} 2\lambda_H v^2 & 2\lambda_{H\sigma} vv_\sigma \\ 2\lambda_{H\sigma} vv_\sigma & 2\lambda_\sigma v_\sigma^2 \end{pmatrix},$$

which has eigenvalues

$$m_H^2 \approx 2 \left( \lambda_H v^2 + \lambda_{H\sigma} v_\sigma^2 \right) \quad \text{and} \quad m_\sigma^2 \approx 2 \left( \lambda_\sigma v_\sigma^2 + \lambda_{H\sigma} v^2 \right).$$

### 2. Theory

#### 2.1. Threshold correction

Consider an energy scale below $m_\sigma < \Lambda_{IS}$, where the heavy scalar $\sigma$ is integrated out. The low-energy Higgs potential should match the SM Higgs potential:

$$V(H) = \lambda_H^{SM} \left( H^\dagger H - \frac{v^2}{2} \right)^2.$$ 

It turns out that the quartic coupling we measure has an additional term:

$$\lambda_H^{SM} = \lambda_H - \lambda_{H\sigma}^2 / \lambda_\sigma.$$ 

Since the SM quartic coupling will be approximately $-0.08$ at $M_P$, the threshold correction

$$\delta \equiv \frac{\lambda_{H\sigma}^2}{\lambda_\sigma}$$

should be large enough to push the high-energy counterpart $\lambda_H$ to positive value all the way up to $M_P$. In the literature there are two possible ways of implementing this threshold mechanism.
(i) One may start by solving the SM RGE’s up to \( m_{\sigma} \), where the SMASH effects kick in, and the quartic coupling \( \lambda_H \) gains a sudden increment by \( \delta \). Continuation of RGE analysis then requires utilizing SMASH RGE’s up to the Planck scale.

(ii) Another way is to solve the SMASH RGE’s from the SM scale, not bothering to solve the low-energy SM RGE’s at all.

We have chosen both approaches respectively in Fig. 1 and 2.

3. Methods

The authors of [7] have generated the two-loop \( \beta \) functions of SMASH. We solve numerically the 14 coupled renormalization group differential equations with respect to Yukawa (\( y_t, y_b, y_{\tau}, y_T, y_N, Y_Q \)), gauge (\( g_1, g_2, g_3 \)) and scalar couplings (\( \mu^2_H, \mu^2_S, \lambda_H, \lambda_\sigma, \lambda_{H\sigma} \)), ignoring the light SM degrees of freedom, from \( M_Z \) to Planck scale. We use MATLAB’s ode45-solver. See Table 1 for used SMASH benchmark points, and Table 2 for our SM input. Our scale convention is \( t \equiv \log_{10}(\mu/\text{GeV}) \).

In some papers, the running of SM parameters (\( y_t, y_b, y_{\tau}, g_1, g_2, g_3, \mu^2_H, \lambda_H \)) obeys the SM RGE’s without corrections from SMASH until \( \mu = \sqrt{2}\mu_S \), after which \( \lambda_H \) gains a threshold correction and the running of all SM parameters follows the SMASH RGE’s. We choose a different approach, where we utilize the SMASH RGE’s from the beginning to all of the parameters.

We scanned over the parameter space \( m_t \in [164, 182] \text{ GeV} \) and \( m_h \in [110, 140] \text{ GeV} \). For every point, we investigated the running of the quartic couplings of the scalar potential. If either \( \lambda_H \) or \( \lambda_\sigma \) turn negative, we denote this point metastable in Fig. 3. If any of the quartic couplings turn negative, we denote this point metastable in Fig. 3.

**Figure 1.** Running of Higgs, \( \sigma \) bare mass and scalar potential parameters with benchmark point. Threshold applied from the beginning at \( m_Z \).
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$1000m_N = v_\sigma = 1e+09 \text{ GeV}, \lambda_S = 2e-10, \lambda_{HS} = -4.7e-06, Y_F = 0.001, Y_Q = 0.001$

**Figure 2.** Running of Higgs, $\sigma$ bare mass and scalar potential parameters with benchmark point. Threshold correction utilized at $m_\rho$.

**Table 1.** Used benchmark points (BPs) in our analysis. Note that we assume specific texture to RH and LH neutrino Yukawa matrices $Y_N = y_N \times \text{diag}(1, 2, 2.1)$ and $Y_F = y_F \times I_3$, respectively, where $y_N$ and $y_F$ are numbers.

<table>
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<th>BP2</th>
<th>BP3</th>
<th>BP4</th>
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<td>$10^{-3}$</td>
<td>$3 \times 10^{-3}$</td>
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<td>$\sqrt{2}/1000$</td>
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<tr>
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<td>$10^{-3}$</td>
<td>$10^{-3}$</td>
</tr>
<tr>
<td>$v_\sigma$</td>
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<td>$10^8 \text{ GeV}$</td>
<td>$10^9 \text{ GeV}$</td>
<td>$10^9 \text{ GeV}$</td>
</tr>
<tr>
<td>$\lambda_{\sigma}$</td>
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<td>$2 \times 10^{-9}$</td>
<td>$10^{-10}$</td>
<td>$1.5 \times 10^{-10}$</td>
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<td>$-10^{-5}$</td>
<td>$-3 \times 10^{-6}$</td>
<td>$-3 \times 10^{-6}$</td>
</tr>
<tr>
<td>$q$</td>
<td>$-1/3$</td>
<td>$-1/3$</td>
<td>$-1/3$</td>
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</tr>
</tbody>
</table>

rises above unity, we denote this point non-perturbative. Analytically a small value of $\lambda_{HS}$ can give positive correction at the one-loop level to push $\lambda_H$ out of the valley of instability, and the correlations of other SMASH parameters to $\lambda_H$ are small.

### 3.1. Choice of benchmark points

To avoid the overproduction of dark radiation via the cosmic axion background, we choose $\lambda_{HS} < 0$. To obtain the observed matter-antimatter asymmetry via leptogenesis, a hierarchy
Table 2. Used SM inputs in our analysis, at $\mu = m_Z = 91.18$ GeV. We define $\mu_H = m_h/\sqrt{2}$ and $\lambda_H = m_h^2/2v^2$.

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<td>$v$</td>
<td>246.22 GeV</td>
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<tr>
<td>$m_b$</td>
<td>4.18 GeV</td>
<td>$g_1$</td>
<td>0.357</td>
</tr>
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<td>$m_{h}$</td>
<td>125.18 GeV</td>
<td>$g_2$</td>
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<td>$m_{\tau}$</td>
<td>1.777 GeV</td>
<td>$g_3$</td>
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</table>

on heavy Majorana neutrinos $N_i$ is required. This is achieved by assigning the right-handed Majorana Yukawa matrix a texture: $Y_N = y_N \times \text{diag}(1,2,2,1)$. As the light neutrinos are also not mass degenerate, we assign the Dirac Yukawa matrix the following texture: $Y_F = y_F \times \text{diag}(1,1,1)$. Values for $y_N$ and $y_F$ are in Table 1.

3.2. Partial unification
The viable parameter space where the SMASH vacuum is stable turns out to give arise of partial gauge sector unification, where $g_2(M_P) \approx g_3(M_P)$.

3.3. Neutrino mass splitting
Basic version of SMASH utilizes Type-I see-saw mechanism. The mass matrix is

\begin{equation}
M_{\nu} = \begin{pmatrix}
0 & M_D & M_M \\
M_D^T & M_D & M_M^T \\
M_M & M_M & M_M^T
\end{pmatrix} = \frac{1}{\sqrt{2}} \begin{pmatrix}
0 & Y_F^T v \\
Y_{F} & Y_{N} v_{\sigma}
\end{pmatrix}
\end{equation}

and

\begin{equation}
m_{\nu} = -M_D M_M^{-1} M_D^T \approx -4 \times 10^8 \text{ GeV} \times \frac{Y_F Y_N^{-1} Y_F^T v_{\sigma}^2}{10^{-6}} \times 10^{-2} \text{ eV}.
\end{equation}

Vanilla leptogenesis scenario requires the existence of heavy RH neutrinos, which are too unstable for the dark-matter candidate. The see-saw scale intermediate between SM and GUT scales, slides well into SMASH framework, with the RH neutrino mass given by VEV of $\sigma$.

Figure 3. Scalar potential vacuum stability regions for $\lambda_{H\sigma} \approx -10^{-5}$ in $(m_H, m_t)$ plane.
Recent results from PLANCK constraints on sum of neutrino masses to be $< 0.12$ eV (95 \% CL.). With the known mass splittings known from oscillation experiments, the heaviest neutrino can be at most $\approx 59$ meV. With a simple Yukawa texture, we are able to fit the present constraints on neutrino masses, see Fig. 4.

![Running of the neutrino masses (mass eigenstates). Mass splitting for atmospheric and solar neutrinos correspond well to present experimental data from oscillation experiments and PLANCK bounds on the sum of light neutrino masses.](image)

4. Conclusions

(i) SMASH unifies axions, see-saw and extended Higgs sector on one energy scale, $\mu \sim 10^{10} - 10^{11}$ GeV, solving several problems badgering the Standard Model in one go.

(ii) SM vacuum is metastable, since $\lambda_H$ turns negative around $\mu \approx 10^{12}$ GeV, SMASH can fix this vacuum metastability problem with $\lambda_{H\sigma} \lesssim 10^{-5}$ at two-loop RGE level.

(iii) Given a suitable neutrino Yukawa texture, SMASH predicts the correct values to atmospheric neutrino mass splitting $\Delta m_{31}^2$ and the solar neutrino mass splitting $\Delta m_{21}^2$.

References